ASTR 1040: Stars & Galaxies

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Lecture 9 Tues 14 Feb 2017

zeus.colorado.edu/astr1040-toomre

Topics for Today and Thur

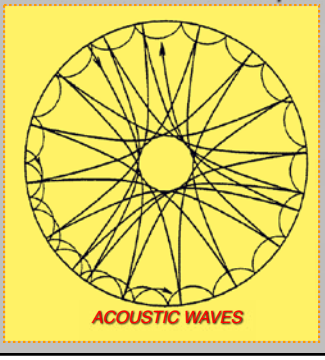
- Helioseismology: acoustic waves excited by convection to probe interior
- Revisit solar magnetism and its cycles
- Use of supercomputers to simulate dynamics within the Sun
- Effects of solar magnetism on Earth

- What can we measure in other stars?
- How do we begin to classify other stars?
- Why temperature and spectral lines are closely linked in classifying stars O B A…M

Logistics

- Overview read Chap 15: Surveying the Stars
- Review Session Wed (tomorrow) 5-7pm here (G130) –- Piyush Agrawal
- Mid-Term Exam 1 Thurs in class (see rules in Review Set, still available)
- Homework #3 (+answers) returned today
- Sorry about Observ #2 (last Thur) cancel; please complete reports from Observ #1

Helioseismology: Millions of sound waves available to probe solar interior



Some waves noodle just below the surface

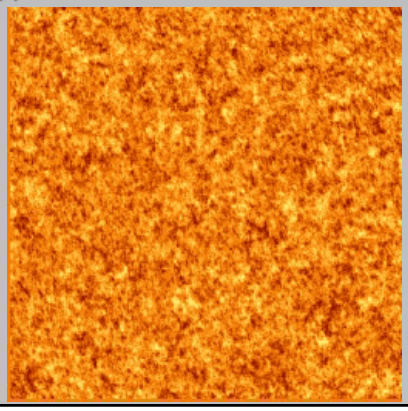
Others almost make it to the center

All excited by turbulent granulation visible in photosphere

How Sound Makes A Surface Bounce

- Sound waves inside Sun cause the photosphere to move up and down, with "five-minute oscillations"
- Can detect these with Doppler imaging of gas at solar surface ("see" the sound)

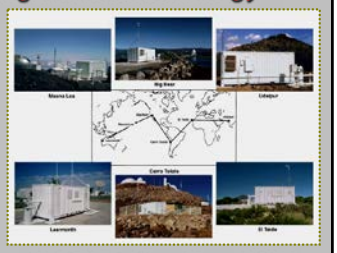
Doppler movie of solar surface from SOHO



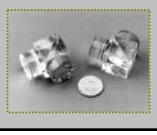
20° across

Tools of Imaging Helioseismology

SOHO Spacecraft
Michelson Doppler Imager (MDI)

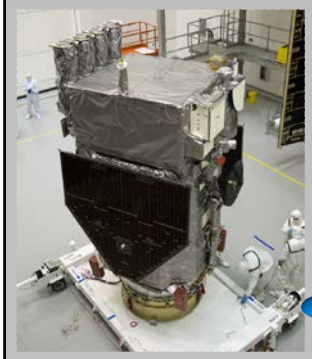


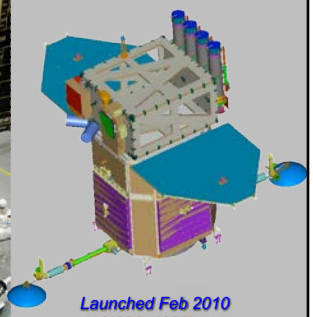
Global Oscillation Network Group (GONG)



Solar Dynamics Observatory (SDO)

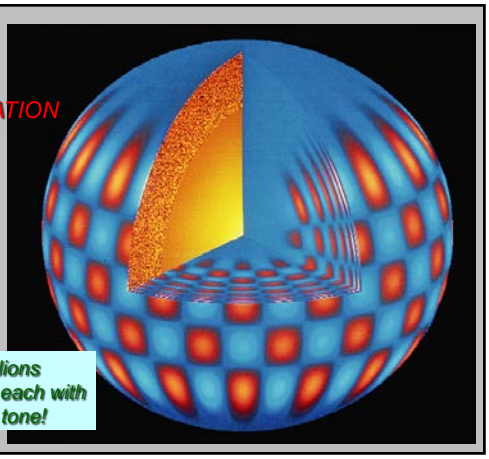
MDI offspring: Helioseismic & Magnetic Imager (HMI)





Launched Feb 2010
(4096x4096)

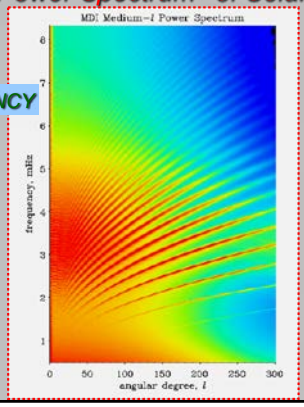
SOLAR OSCILLATION MODE



One of millions of modes, each with a different tone!

"Power Spectrum" of Solar Oscillations

FREQUENCY



Each oscillation mode has a distinctive (measurable) tone

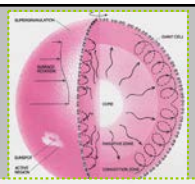
Sensitive to how sound speed varies with depth

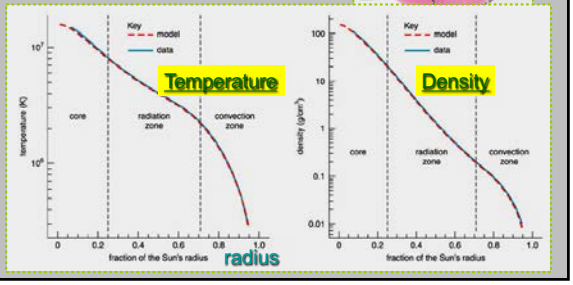
Tests models of inside temperature

Observed from SOHO and GONG

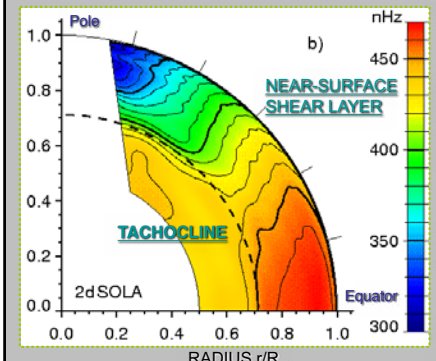
Solar Interior Models

Theory vs Helioseismology





Interior Differential Rotation Profile



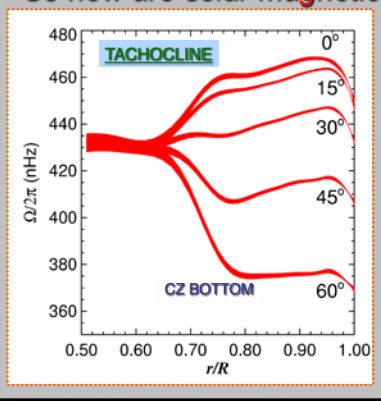
Big surprises:

Fast equator, slow pole -- over most of convection zone

Tacholine of strong shear at its base

Radiative zone rotates uniformly

So how are solar magnetic fields built?

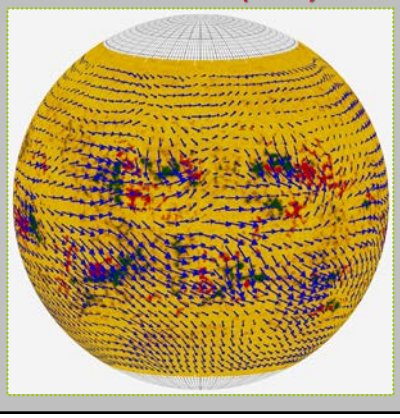


TWO MAGNETIC DYNAMOS:

Global magnetic fields built near TACHOCLINE

Small-scale fields built in near-surface shear layer

Solar Subsurface Weather (SSW) Global View

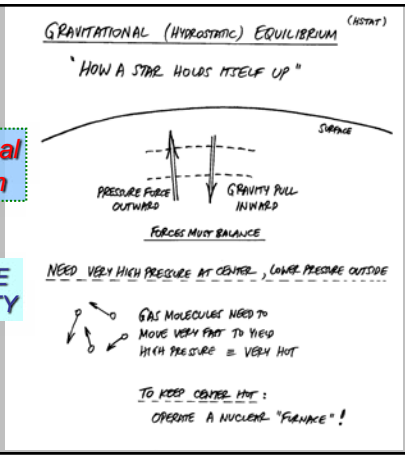


April 2002

REVISIT

Gravitational equilibrium

PRESSURE vs GRAVITY



What can a pressure gradient do for you?

Hold up a star, crush a steel drum ….

Reasoning Clicker Q

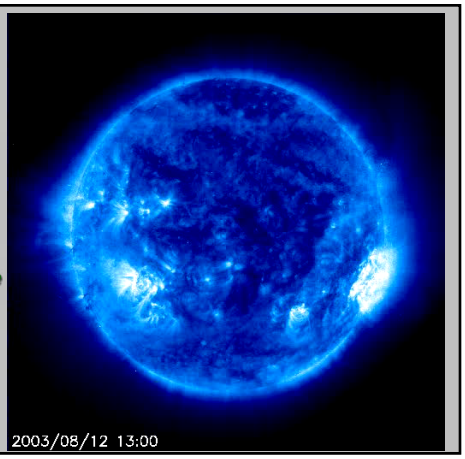
B.

- If the Sun's core went out of balance and shrank a little, what would happen there?
- A. Density would decrease and fusion would slow down, releasing less energy
- B. Density and temperature would increase and fusion would speed up, releasing more energy
- C. The whole Sun would eventually shrink and thus core would come back into balance
- D. Not much would really change, so nothing to worry about

RETURN TO MAGNETISM

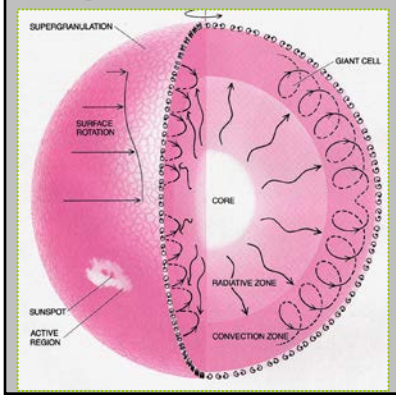
UV view of solar magnetism

chromosphere & low corona



SOHO EIT 171 A

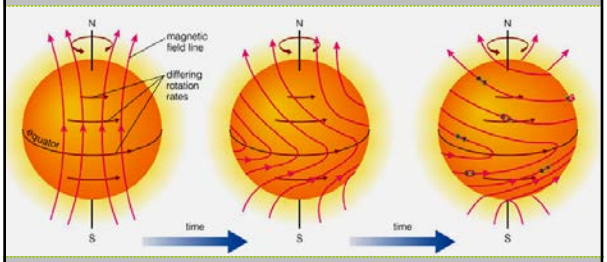
Magnetic fields are built in convection zone

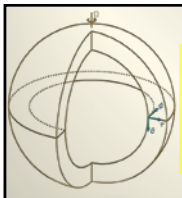


Convection Zone:
ROTATING SHELL OF TURBULENT PLASMA

Differential Rotation and Magnetic Dynamo
.....
go HAND-IN-HAND

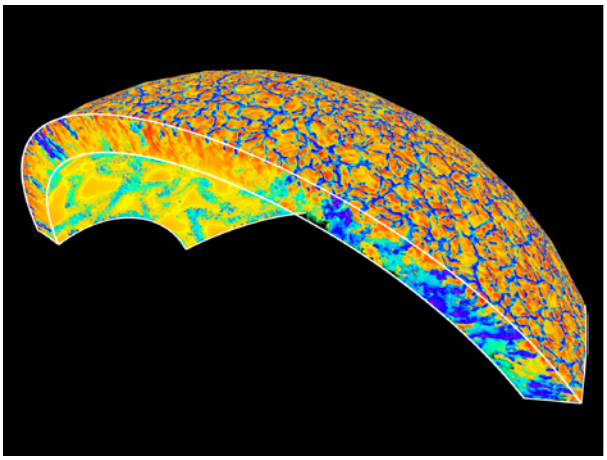
How differential rotation can stretch and change global magnetic fields



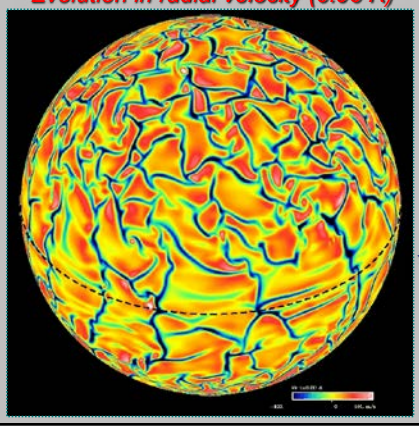


Supercomputer simulations:
Global convection & dynamo building of magnetic fields

- Study turbulent convection interacting with rotation in bulk of solar CZ: 0.72 R - 0.97 R
- Realistic solar mean stratification
- Simplified physics: perfect gas, radiative diffusivity, compressible, subgrid transport
- Correct global spherical geometry



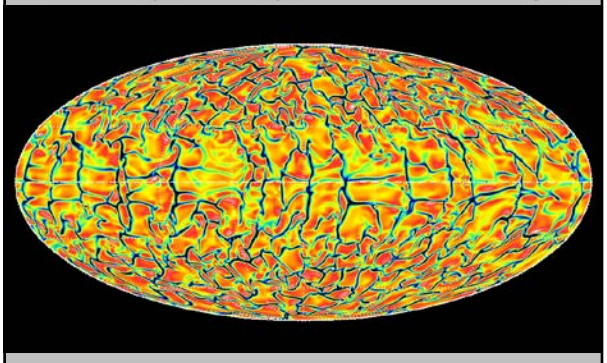
Evolution in radial velocity (0.96 R)



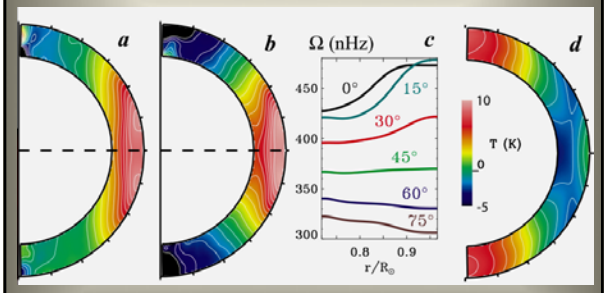
NEAR TOP OF LAYER

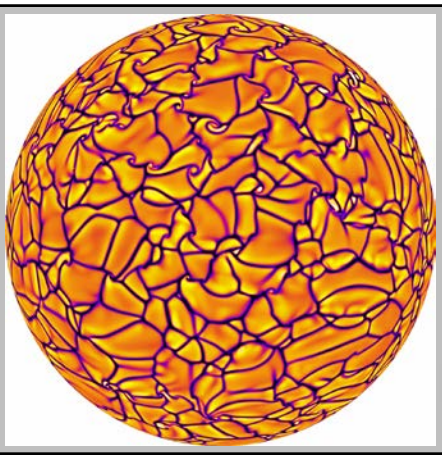
VIEWED IN CARRINGTON FRAME

Global (Mollweide) view of radial velocity



Resulting differential rotation built by convection: fast equator, slower poles





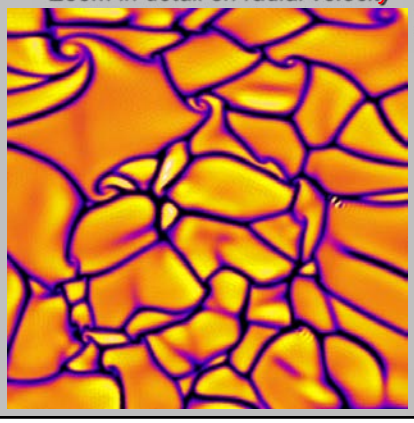
At highest resolution

Radial velocity

Case F

Miesch et al 2008, 2011

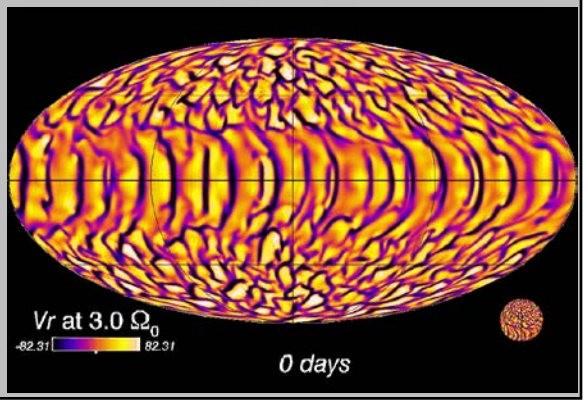
Zoom-in detail on radial velocity

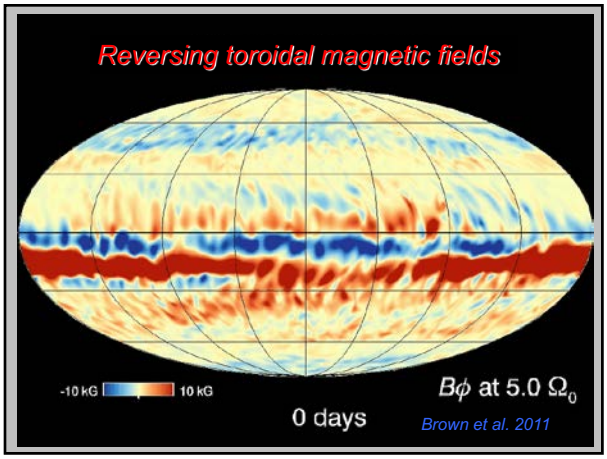


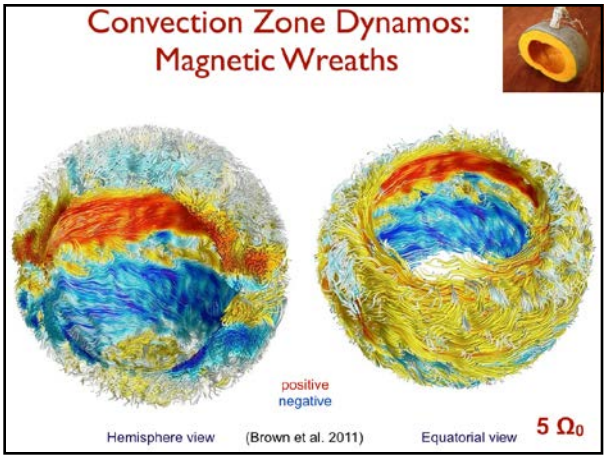
50° N

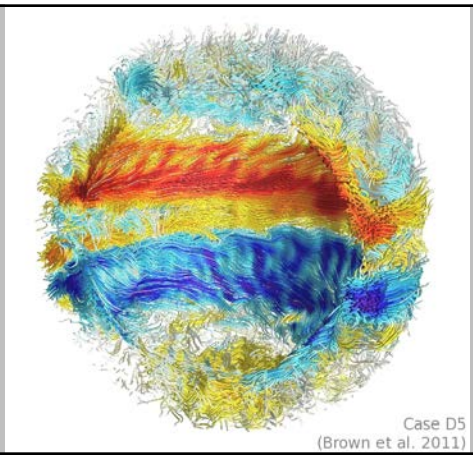
EQ

Now allow MHD dynamo to be turned on

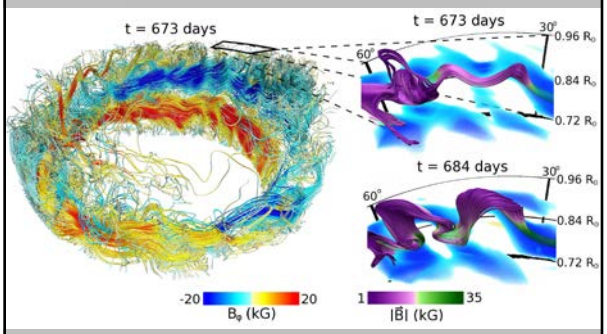








Buoyant magnetic loops ….
may emerge as sunspot loops

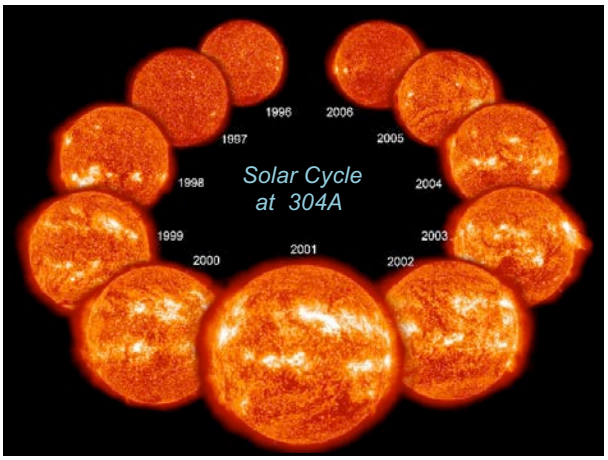


Reading Clicker – Solar Maximum ?

- *What observed features characterize the Sun at "solar maximum"?*

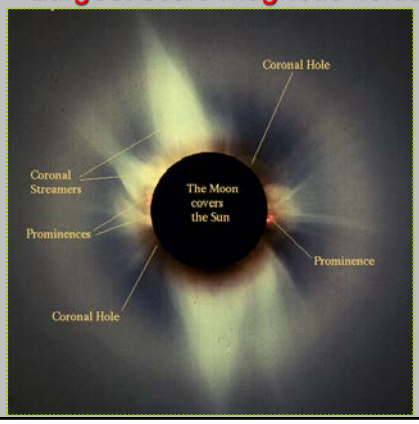
D.

- **A.** Sun becomes much brighter
- **B.** Sun emits light of longer wavelengths
- **C.** Sun rotates faster at the equator
- **D.** Many sunspots are visible on surface
- **E.** All of the above



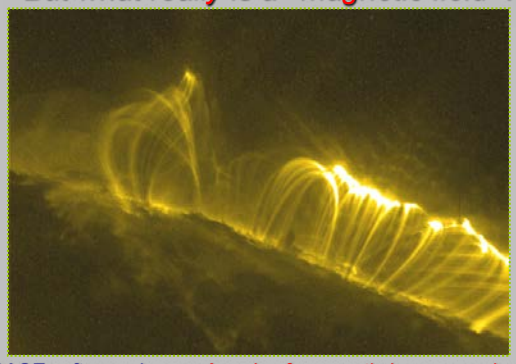
EUV from SOHO

Largest-scale magnetic fields



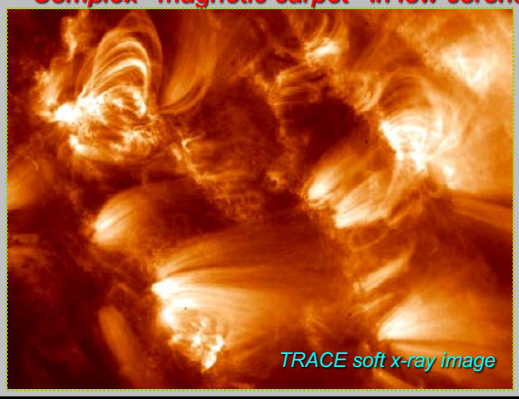
Sunspots at surface yields huge magnetic structures in corona

But what really is a "magnetic field"?



TRACE soft x-ray image: Arcade of magnetic loops on solar limb

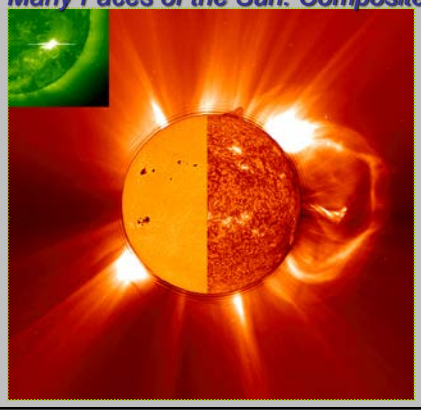
Complex "magnetic carpet" in low corona



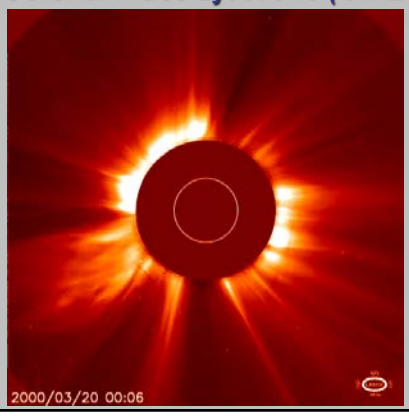
Magnetic Reconnection and Splendid Loops from SDO

Huge prominence is big magnetic loop

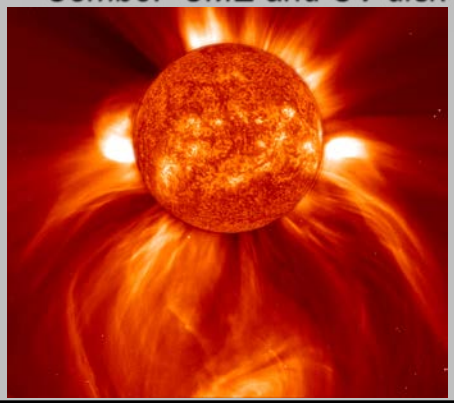
Many Faces of the Sun: Composite



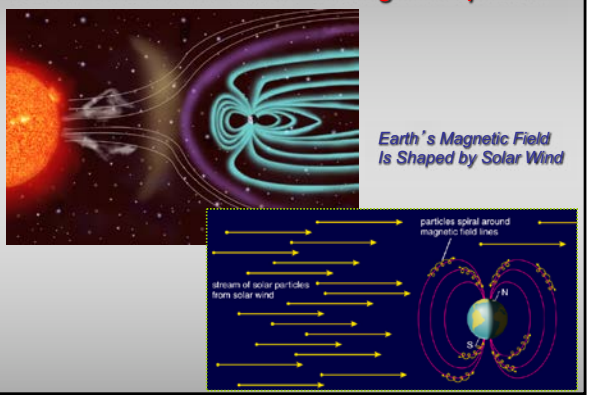
Coronal Mass Ejections (CMEs)



Combo: CME and UV disk



Solar Wind and Earth's Magnetosphere



Solar Wind and Aurorae

Northern Lights
(Aurora Borealis)

Reading Clicker Question

Which is the most likely cause of the extreme heating in the chromosphere and corona?

A. *Energy deposited by magnetic fields*
B. *Heat rising from the surface of the Sun*
C. *Photons created at the photosphere interacting with the solar atmosphere*
D. *Neutrino interactions with the solar wind*
E. *Ionization of hydrogen just above the surface*